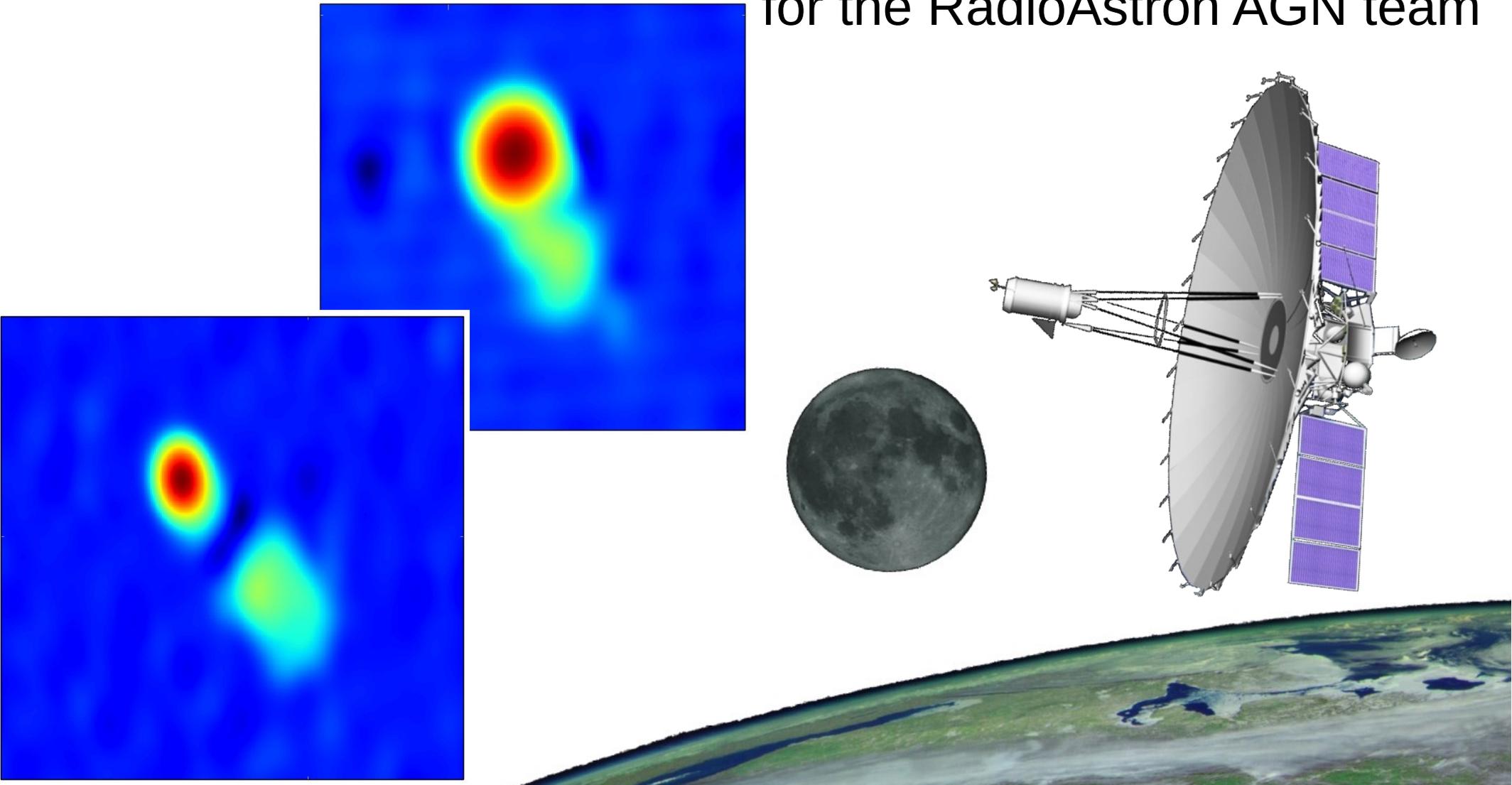


# RadioAstron 5-22 GHz observations of 3C 418 and 2013+370

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for the RadioAstron AGN team



# Aims

## *Technical test*

- Check amplitude calibration of the Space radio telescope through hybrid imaging
- Test RadioAstron's dual-band imaging capability at 4.8 and 22 GHz (simultaneously)
- Test RadioAstron's ability to image two nearby ( $15^\circ$  apart) sources in one experiment

## *High-resolution imaging*

- Determine *shape*, *size* and  $T_b$  of the most compact jet structures (optically-thick cores)

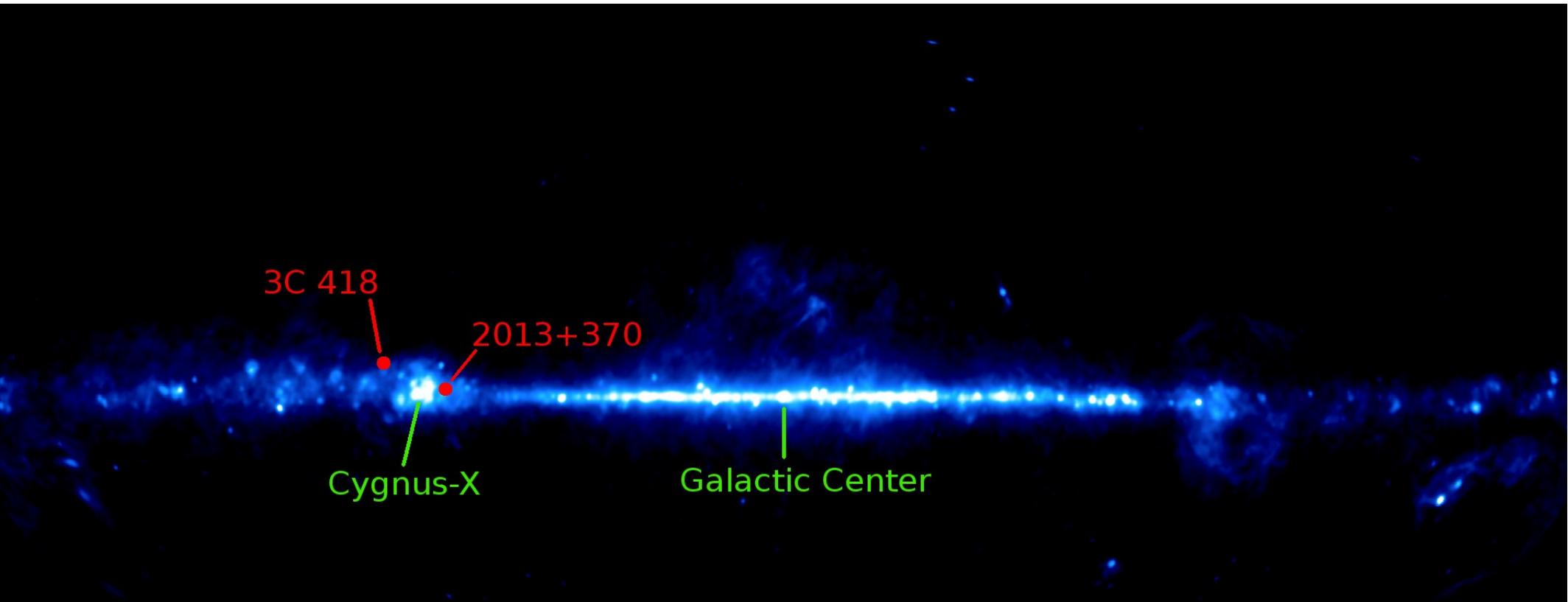
# Target selection

Bright VLBI sources with space-ground baselines crossing ground-ground baselines in October 2012

- **TXS 2013+370** FSRQ at  $z=0.859$ , GeV-bright, 3mm VLBI detected (Lee et al. 2008, AJ, 136, 159),  $\beta=12.5c$  (MOJAVE)
- **3C 418** FSRQ at  $z=1.686$ , no GeV detection, 3mm VLBI detected,  $\beta=3.75c$

# 3C 418 and TXS 2013+370

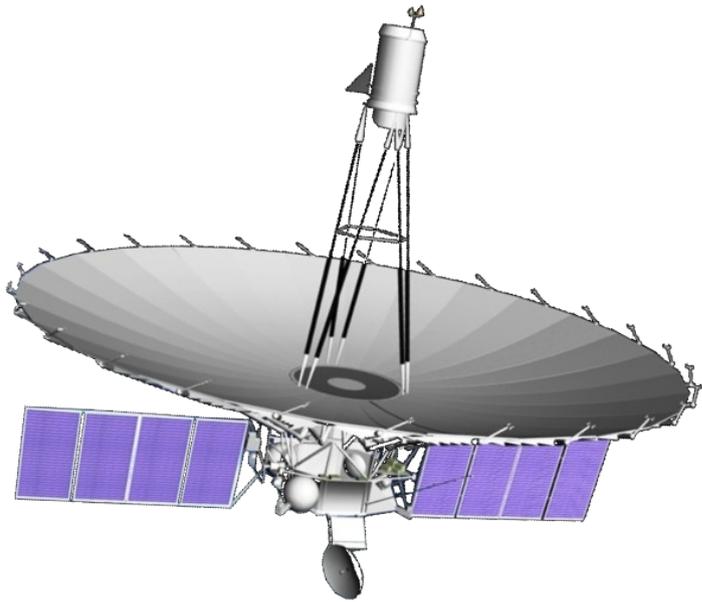
are close to the Galactic plane at  $b=6.0^\circ$  and  $b=1.2^\circ$ .  
Expect interstellar scattering.



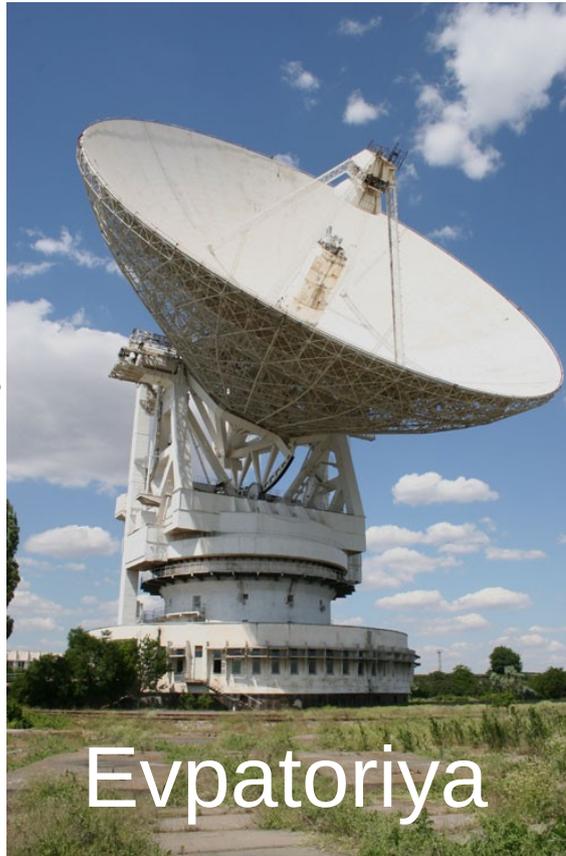
WMAP 9-year all-sky 23 GHz map  
Bennett et al. 2013, ApJS, 208, 20

# Array configuration

EVN + Usuda 64m + Evpatoria 70m divided in two subarrays observing at 4.8 and 22 GHz, the Space telescope observing in two bands simultaneously



Space radio  
telescope



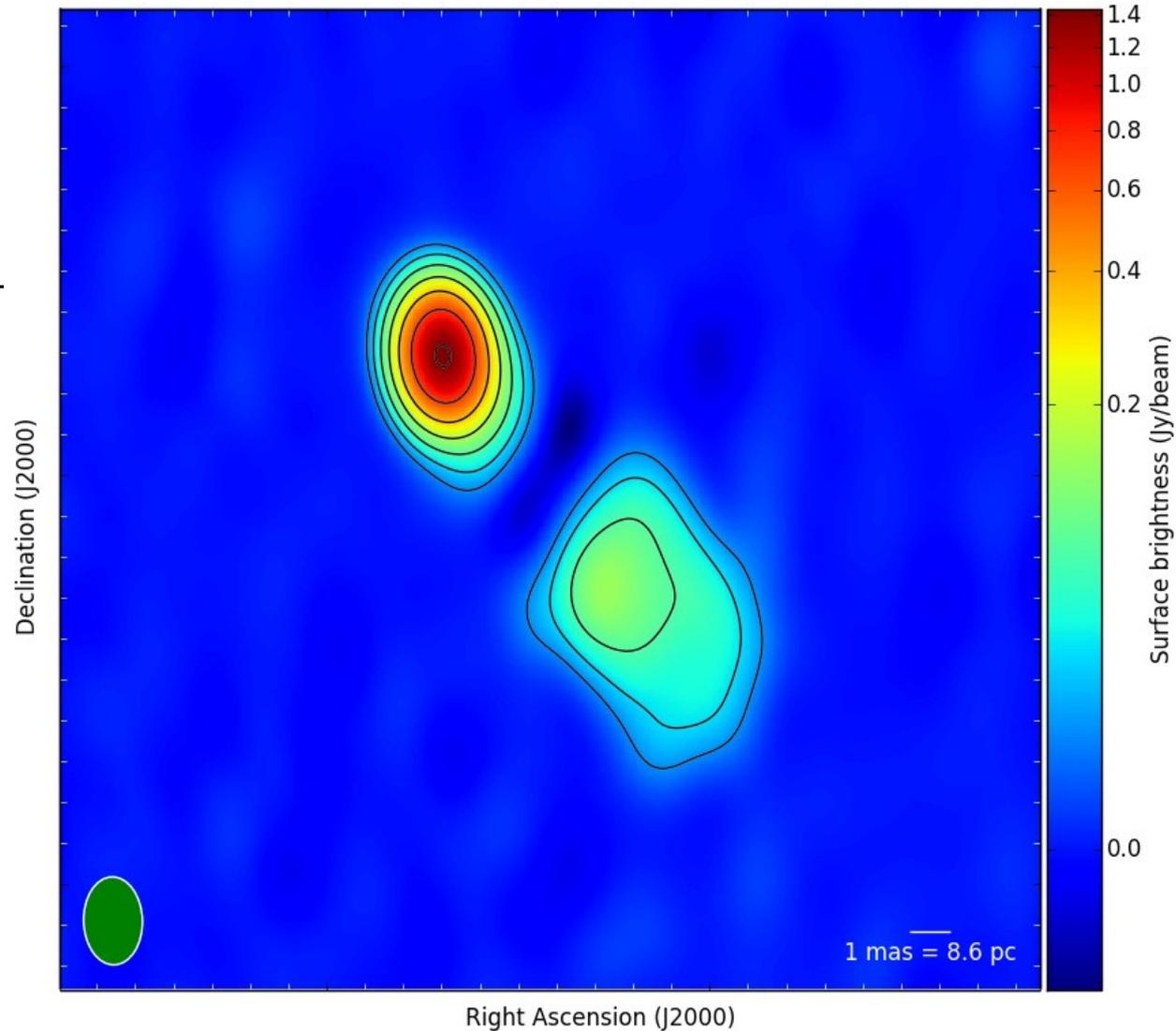
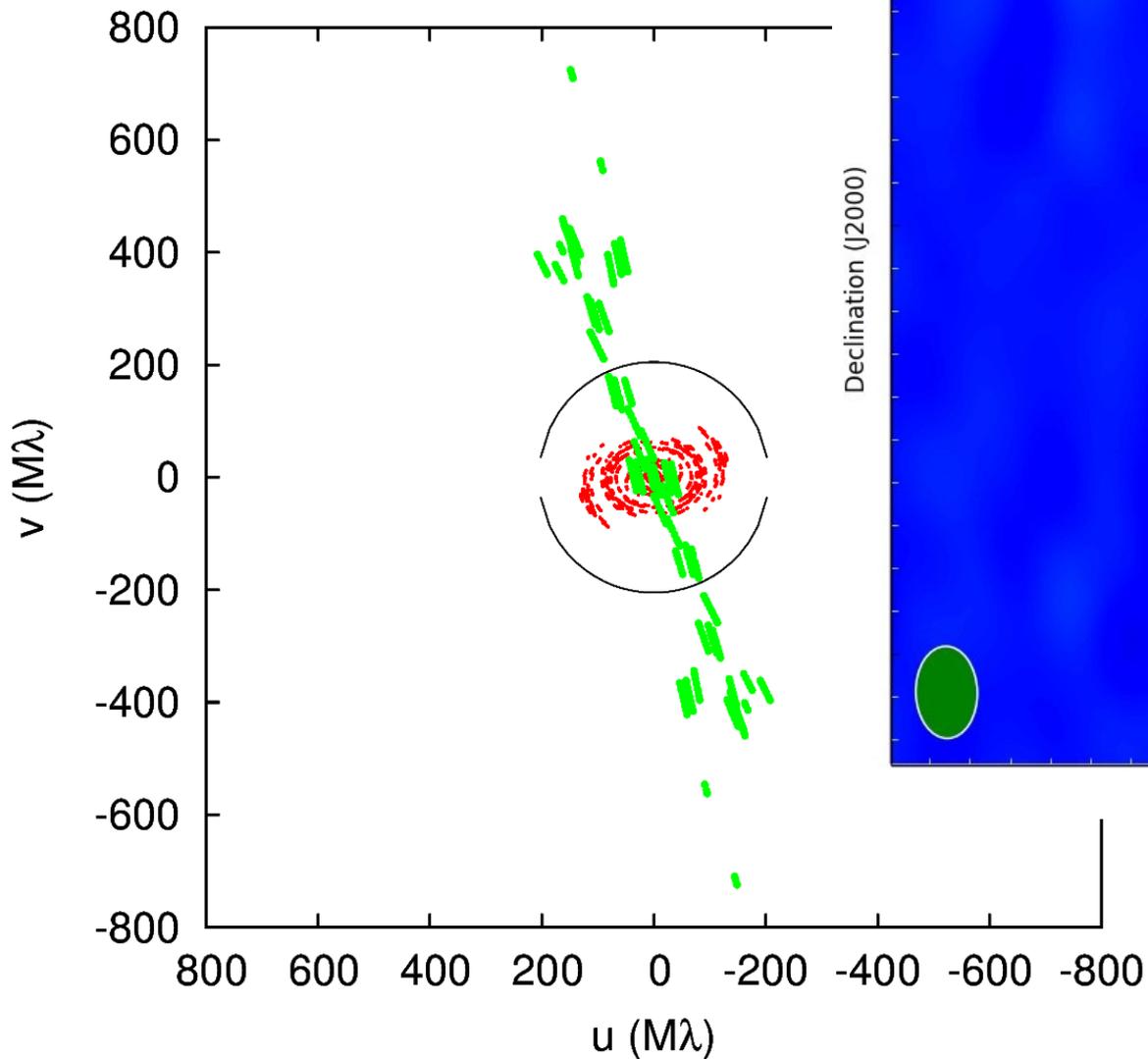
Evpatoriya



WSRT

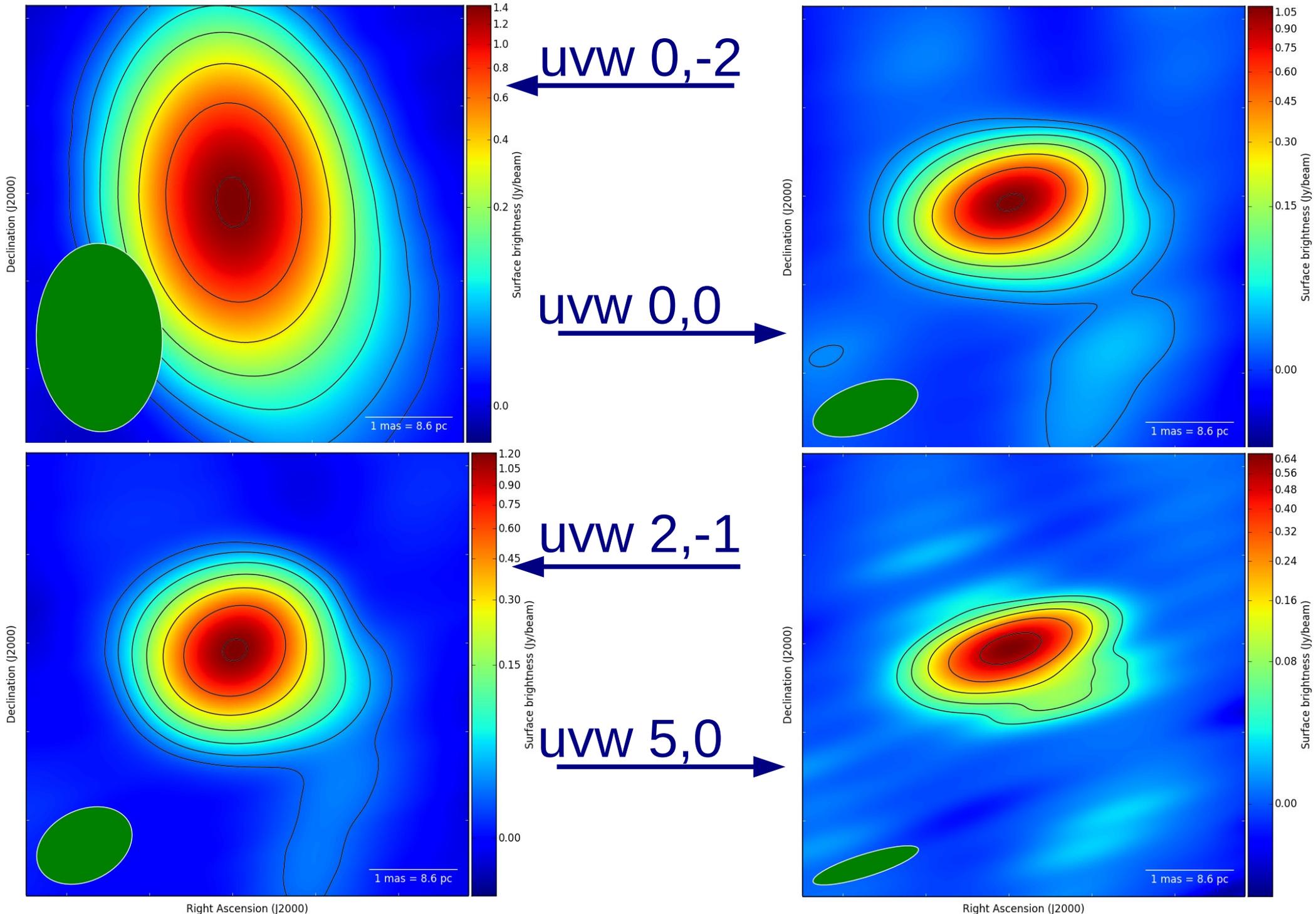
# 3C 418 at 4.8 GHz

uv-coverage

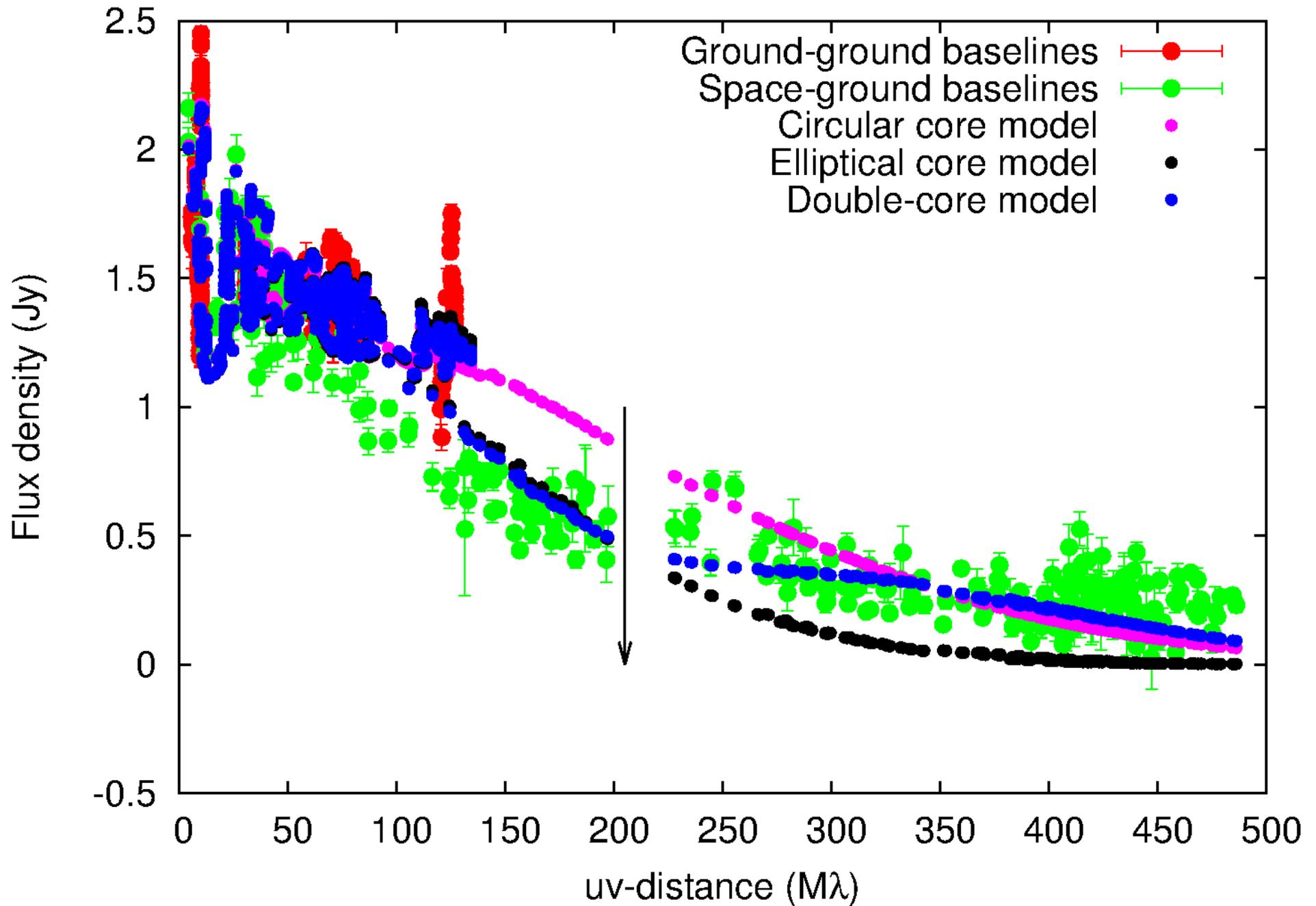


naturally-weighted  
image (uvw 0,-2)

# 3C 418 at 4.8 GHz



# 3C 418 at 4.8 GHz

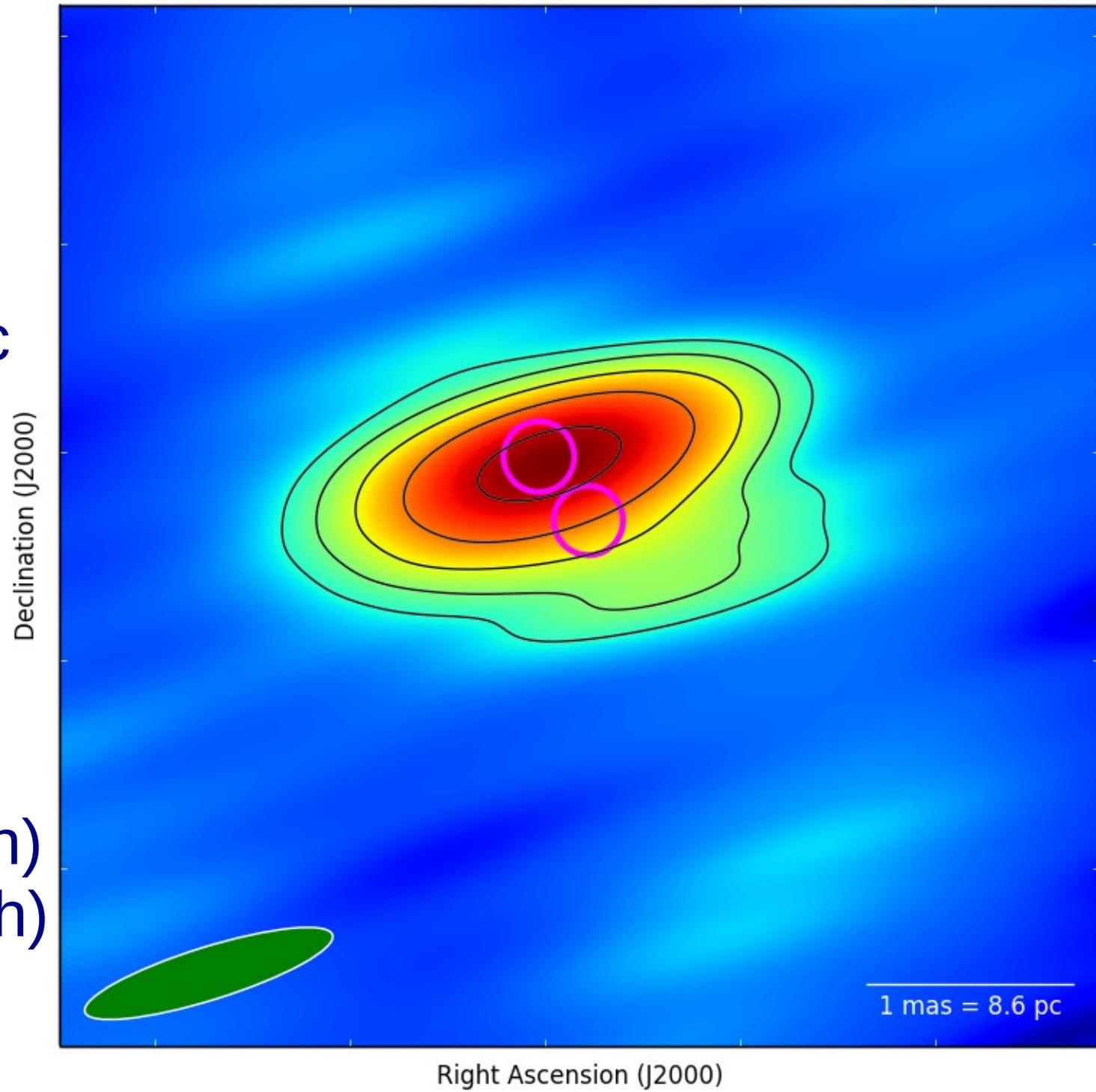


# Two-component model of 3C 418's core

Size  
0.3 mas = 2.6 pc  
(each)

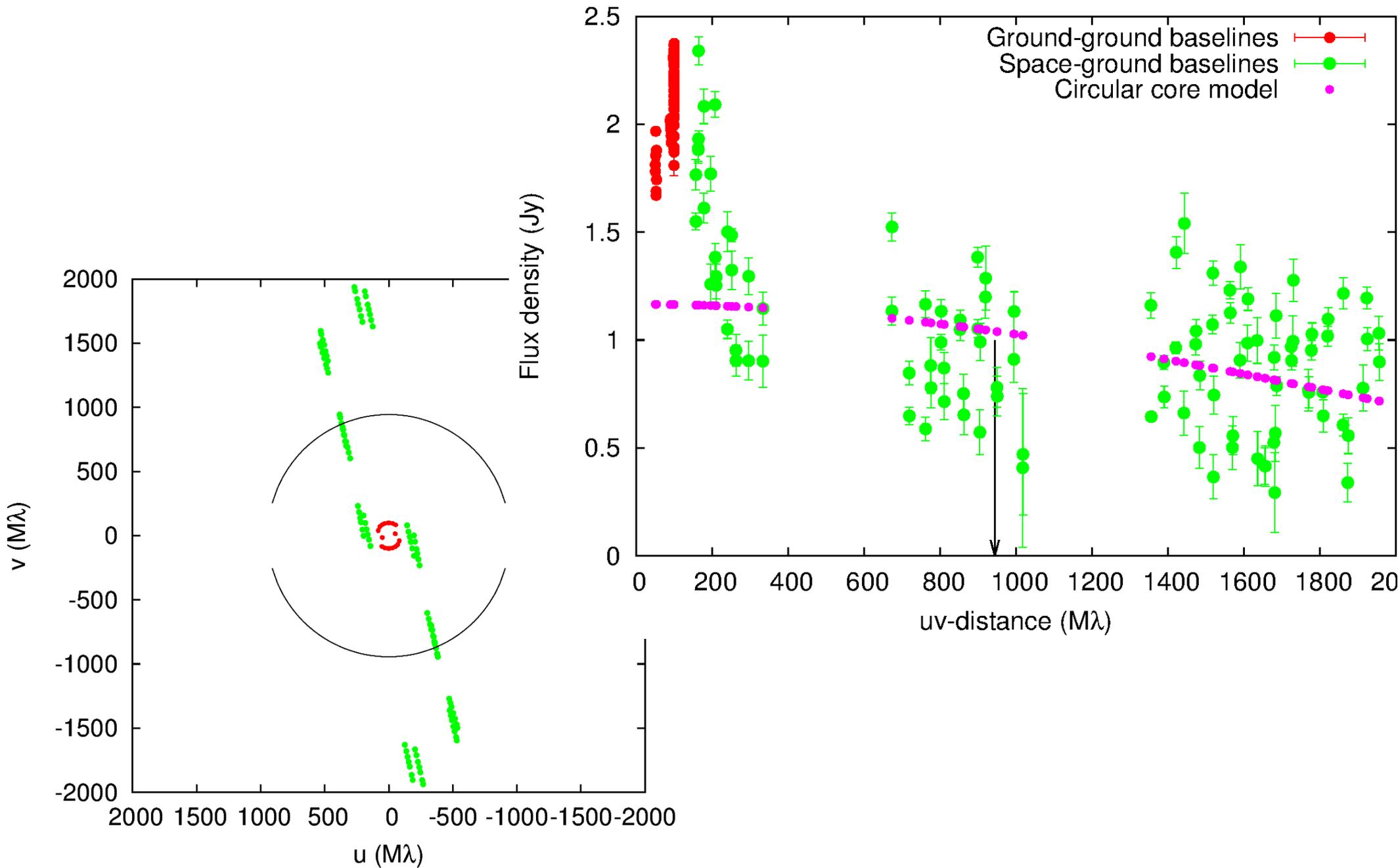
Flux densities  
1.14 & 0.40 Jy

Tb  
 $1 \times 10^{12}$  K (north)  
 $4 \times 10^{11}$  K (south)

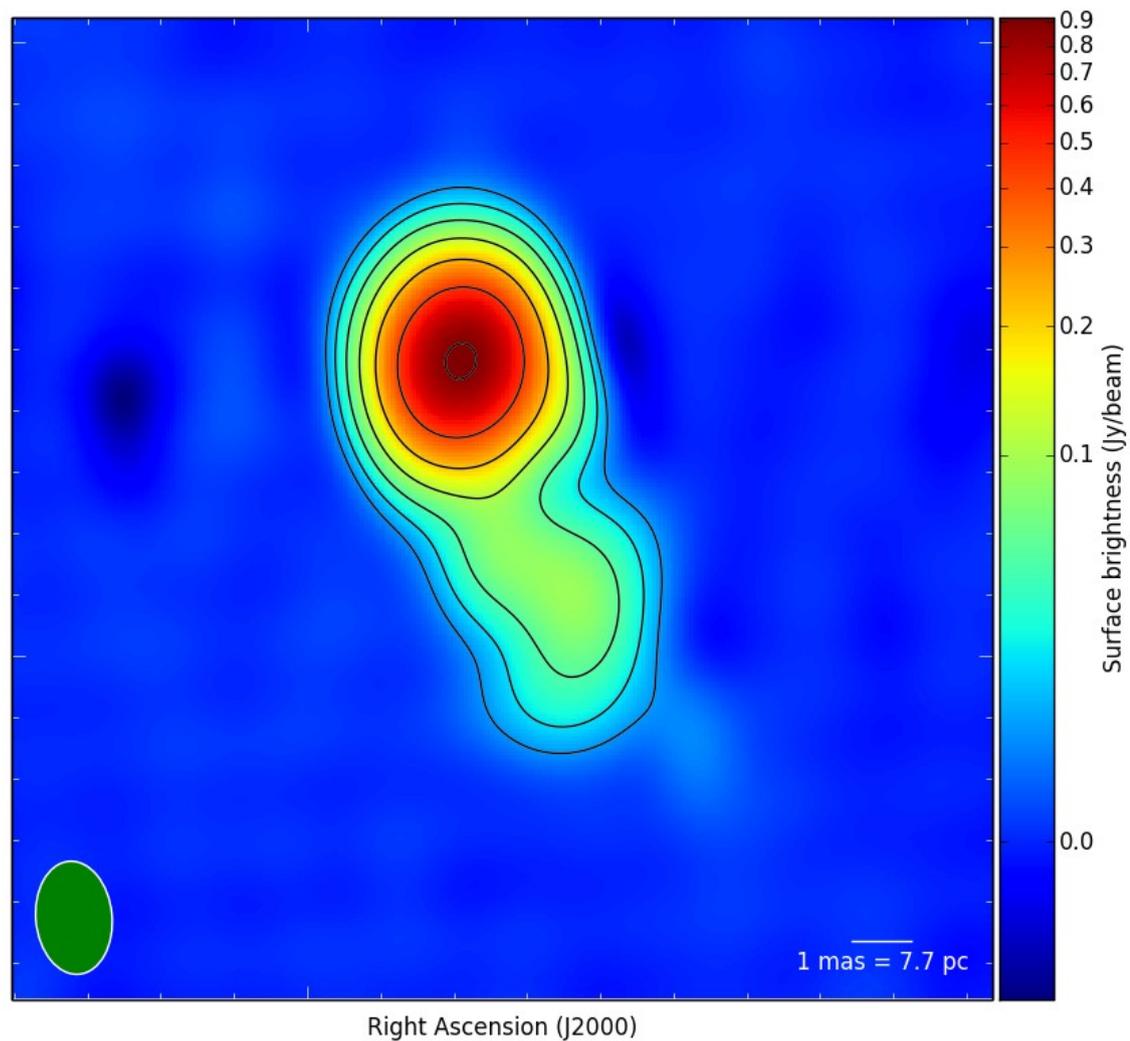
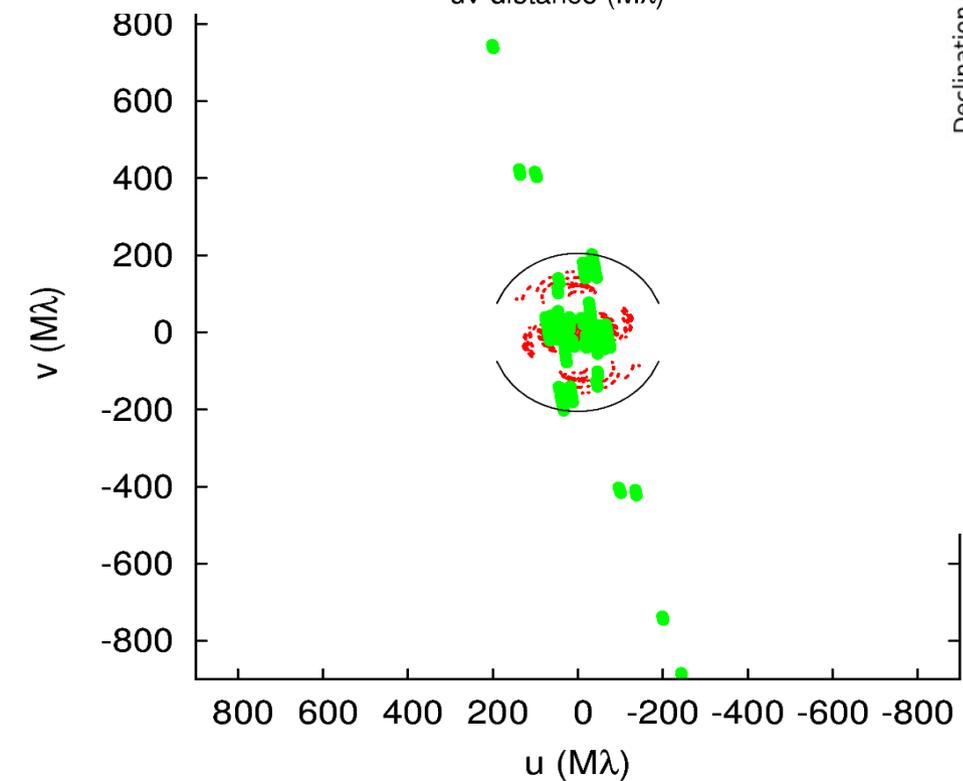
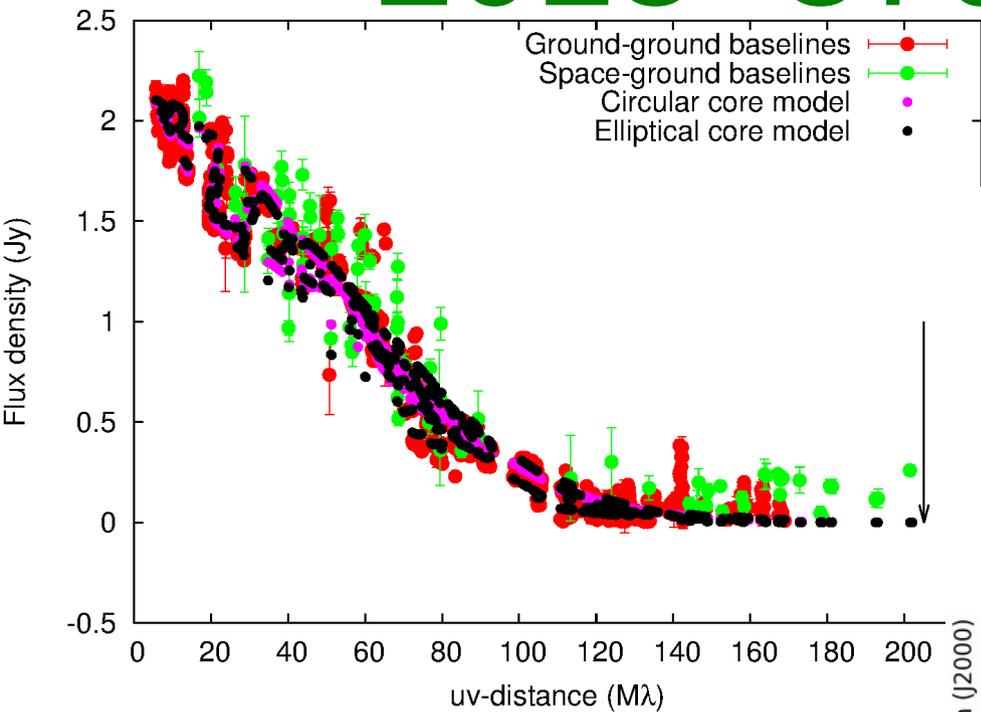


# 3C 418 at 22 GHz

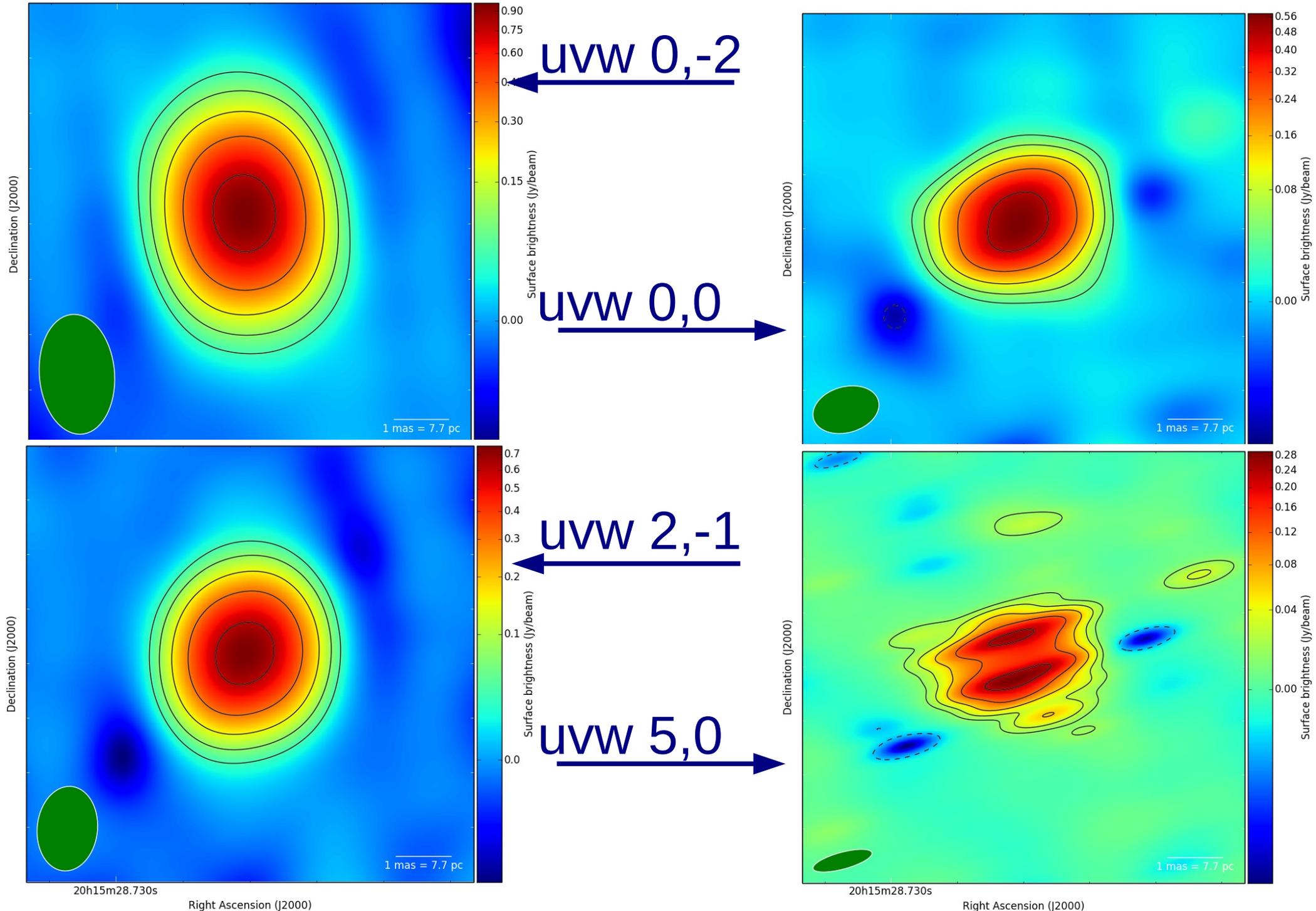
0.039 mas = 0.33 pc, 1.17 Jy,  $T_b = 5 \times 10^{12}$  K



# 2013+370 at 4.8 GHz



# 2013+370 at 4.8 GHz

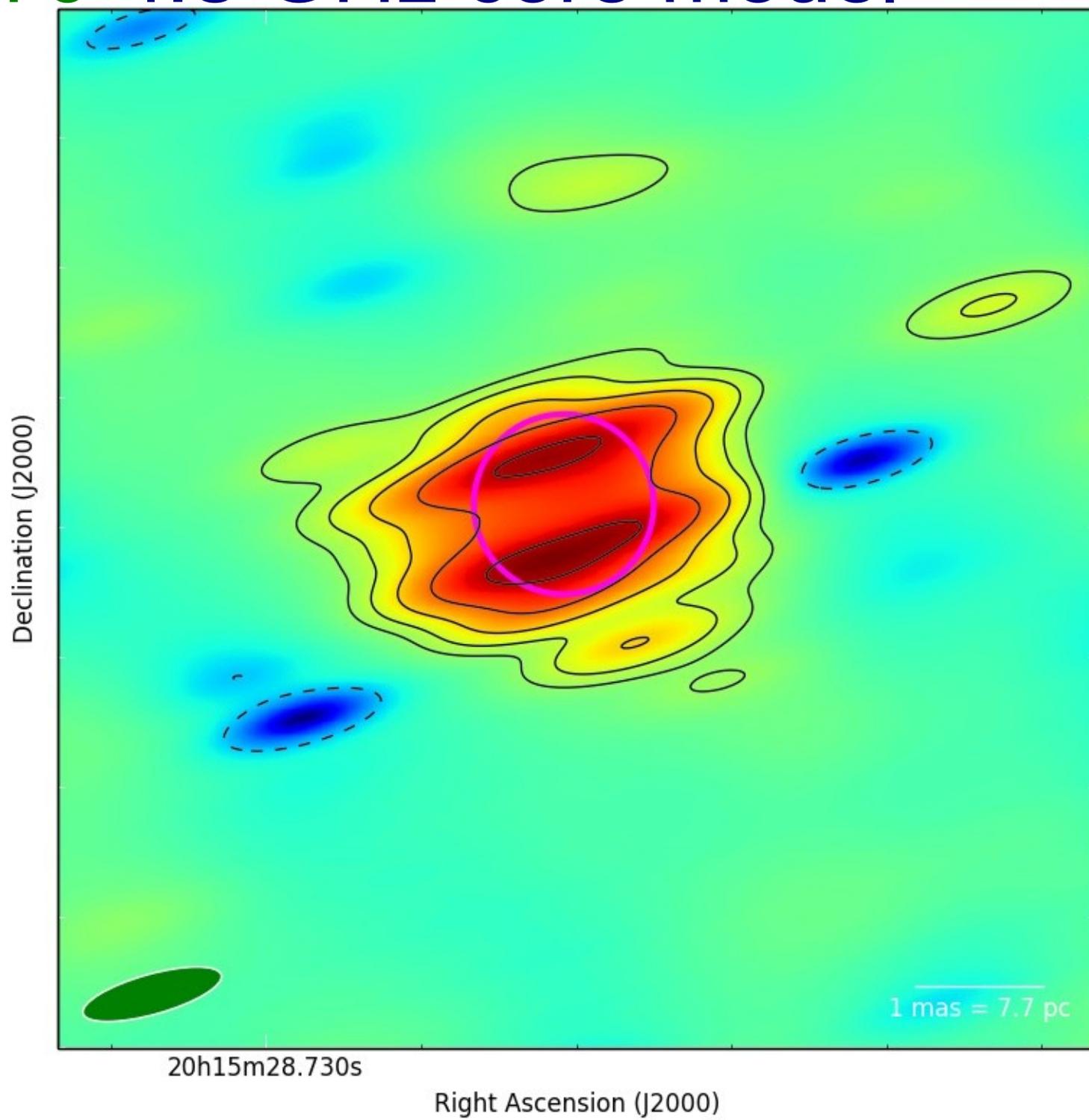


# 2013+370 4.8 GHz core model

Size  
1.5 mas = 12 pc

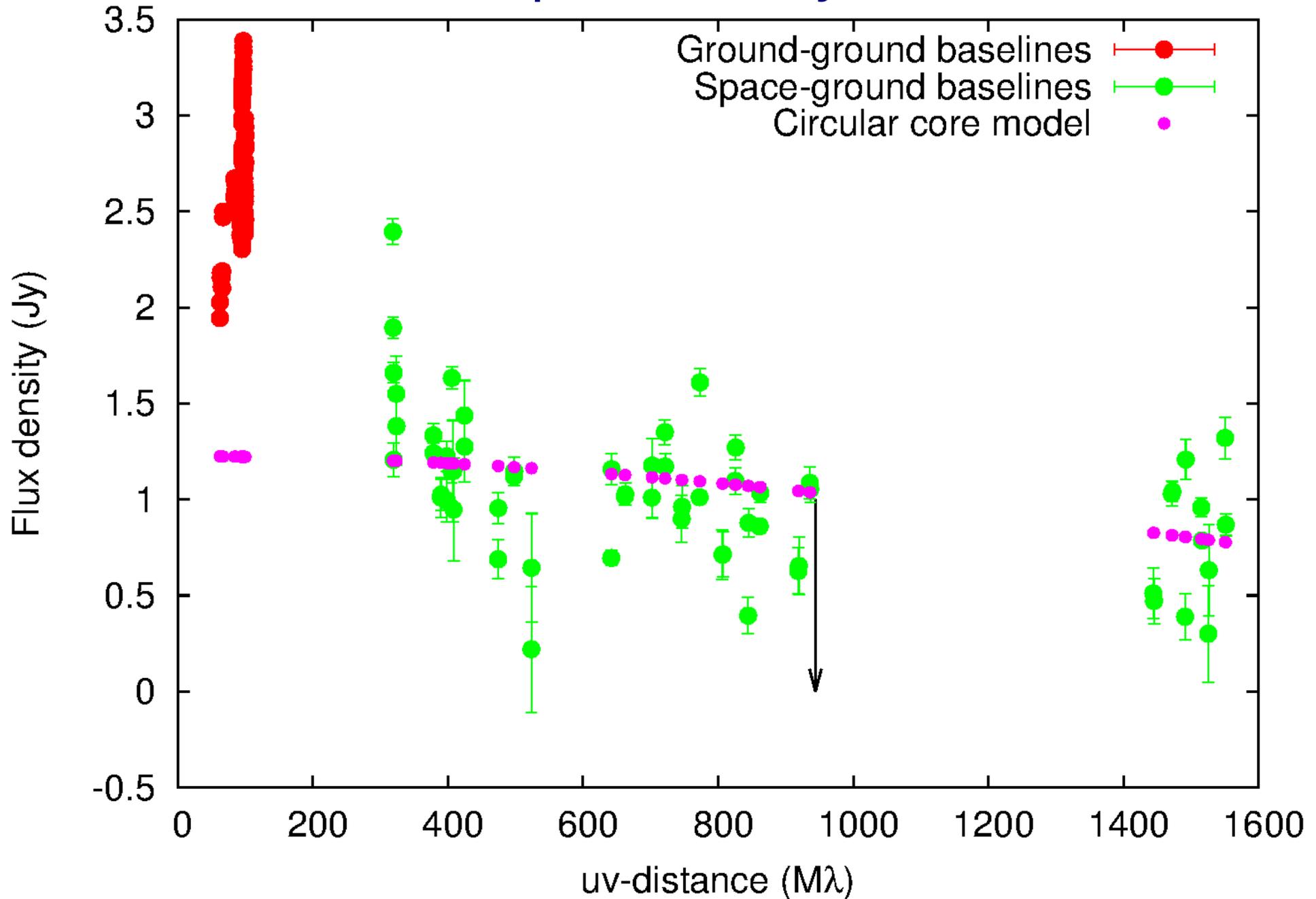
Flux density  
1.86 Jy

Tb  
 $8 \times 10^{10}$  K



# 2013+370 at 22 GHz

0.047 mas = 0.37 pc, 1.22 Jy,  $T_b=2.5 \times 10^{12}$  K



# Conclusions

- Both sources resolved down to tens of mJy level
- Cores of both sources are likely scatter-broadened (4.8 GHz core size 9-30 times larger than 22 GHz size)
- **3C 418** – complex core structure that cannot be recovered with ground-only obs. at 4.8 GHz
- $T_b > 10^{12}$  K found at 22 GHz

Backup slides...

# Structure or scattering?

- If the source size is fully determined by scattering, its size is expected to scale as  $\lambda^2$
- If the scattering is not important, the core size is expected to scale as  $\lambda^1$  (BK-type jet with SSA)
- $6.2\text{cm}/1.35\text{cm} = 4.6$ ,  $(6.2\text{cm}/1.35\text{cm})^2 = 21.1$
- 3C 418:  $0.37\text{mas}/0.039\text{mas} = 9.5$
- TXS 2013+370:  $1.53\text{mas}/0.047\text{mas} = 32$

# Array configuration

Space + 9 (8 for 3C 418 invisible for Hh) ground telescopes collected useful 4.8 GHz data

C-band ground subarray baseline lengths (km)

	Wb	Jb1	On	Tr	Sv	Bd	Ur	Sh	Hh	Ev	Ud
Wb	0	599	601	799	1634	5786	5565	8090	8239	2097	8347
Jb1	599	0	1011	1388	2032	6155	6028	8419	8441	2683	8578
On	601	1011	0	637	1080	5272	5119	7647	8525	1987	7885
Tr	799	1388	637	0	1070	5199	4874	7552	8108	1375	7925
Sv	1634	2032	1080	1070	0	4281	4127	6760	8697	1716	7074
Bd	5786	6155	5272	5199	4281	0	1452	2749	9832	4839	3293
Ur	5565	6028	5119	4874	4127	1452	0	3249	8852	4152	4303
Sh	8090	8419	7647	7552	6760	2749	3249	0	10160	7067	1680
Hh	8239	8441	8525	8108	8697	9832	8852	10160	0	7391	11085
Ev	2097	2683	1987	1375	1716	4839	4152	7067	7391	0	7721
Ud	8347	8578	7885	7925	7074	3293	4303	1680	11085	7721	0

# Array configuration

Space + 3 ground telescopes collected  
useful 22 GHz data

K-band ground subarray baseline lengths (km)

	Ef	Jb2	Ys	Nt	Gb	Ro
Ef	0	699	1352	1644	6335	1413
Jb2	699	0	1411	2247	5719	1427
Ys	1352	1411	0	1616	6124	99
Nt	1644	2247	1616	0	7446	1711
Gb	6335	5719	6124	7446	0	6049
Ro	1413	1427	99	1711	6049	0

# Correlation and post-processing

- RA-enabled DiFX (J. Anderson)
- Preliminary correlation done in MPIfR-Bonn
- Fringe search in PIMA (L. Petrov)
- Final DiFX correlation in ASC (slow but flexible)
- Fringe fitting in PIMA including accel. term (rate-rate)
- Imaging/modeling in Difmap
- TODO: repeat correlation using the ASC correlator and SFXC (JIVE), compare results